

Reynolds numbers in the early Universe

Massimo Giovannini¹

Department of Physics, Theory Division, CERN, 1211 Geneva 23, Switzerland

INFN, Section of Milan-Bicocca, 20126 Milan, Italy

Abstract

After electron-positron annihilation and prior to photon decoupling the magnetic Reynolds number is approximately twenty orders of magnitude larger than its kinetic counterpart which is, in turn, smaller than one. In this globally neutral system the large-scale inhomogeneities are provided by the spatial fluctuations of the scalar curvature. Owing to the analogy with the description of Markovian conducting fluids in the presence of acoustic fluctuations, the evolution equations of a putative magnetic field are averaged over the large-scale flow determined by curvature perturbations. General lessons are drawn on the typical diffusion scale of magnetic inhomogeneities. It is speculated that Reynolds numbers prior to electron-positron annihilation can be related to the entropy contained in the Hubble volume during the various stages of the evolution of the conducting plasma.

¹e-mail address: massimo.giovannini@cern.ch

In a conducting plasma, such as the early Universe, the kinetic and magnetic Reynolds numbers are defined as [1, 2, 3]

$$R_{\text{kin}} = \frac{v_{\text{rms}} L_v}{\nu_{\text{th}}}, \quad R_{\text{magn}} = \frac{v_{\text{rms}} L_B}{\nu_{\text{magn}}}, \quad Pr_{\text{magn}} = \frac{R_{\text{magn}}}{R_{\text{kin}}}, \quad (1)$$

where v_{rms} estimates the bulk velocity of the plasma while ν_{th} and ν_{magn} are the coefficients of thermal and magnetic diffusivity; L_v and L_B are, respectively, the correlation scales of the velocity field and of the magnetic field. In Eq. (1) Pr_{magn} denotes the so-called magnetic Prandtl number [1, 2, 3].

Prior to electron-positron annihilation (i.e. $T \geq \text{MeV}$) the coefficient of thermal diffusivity can be estimated as $\nu_{\text{th}} \sim (\alpha_{\text{em}}^2 T)^{-1}$ from the two-body scattering of relativistic species with significant momentum transfer. The conductivity of the plasma is $\sigma \sim T/\alpha_{\text{em}}$ so that the magnetic diffusivity becomes $\nu_{\text{magn}} = \alpha_{\text{em}}(4\pi T)^{-1}$. Assuming, for sake of simplicity, thermal and kinetic equilibrium of all relativistic species (which is not exactly the case for $T \sim \text{MeV}$) the kinetic Reynolds number turns out to be $R_{\text{kin}} \simeq \mathcal{O}(10^{16})$, the magnetic Reynolds number is $R_{\text{magn}} \simeq 4\pi/\alpha_{\text{em}}^3 R_{\text{kin}} \sim \mathcal{O}(10^{24})$ and $Pr_{\text{magn}} \sim 10^7$. The latter estimates have been obtained by assuming, in Eq. (1), $L_v \simeq L_B \sim H^{-1}$ (where H^{-1} is the Hubble radius at the corresponding epoch); when the evolution of the background geometry is decelerated (i.e. $a(t) \sim t^\epsilon$ with $0 < \epsilon < 1$) = the particle horizon coincides with the Hubble radius up to an immaterial numerical factor (i.e. $\epsilon/(1-\epsilon)$) which shall be neglected throughout². In the symmetric phase of the standard electroweak theory, picking up a temperature $T > 100 \text{ GeV}$ where all the species (including the Higgs boson and the top quark) are in thermal and kinetic equilibrium, ν_{th} and ν_{magn} can be computed [4] in terms of the hypercharge coupling constant since the non-screened vector modes at finite conductivity are associated with the hypercharge field. In the electroweak case $R_{\text{kin}} \sim \mathcal{O}(10^{11})$ and $R_{\text{magn}} \sim \mathcal{O}(10^{17})$.

The hypothesis of primeval turbulence has been a recurrent theme since the first speculations on the origin of the light nuclear elements. The implications of turbulence for galaxy formation have been pointed out in the fifties by Von Weizsäcker and Gamow [5]. They have been scrutinized in the sixties and early seventies by various authors [6] (see also [7, 8] and discussions therein). In the eighties it has been argued [9] that first-order phase transitions in the early Universe, if present, can provide a source of kinetic turbulence and, hopefully, the possibility of inverse cascades which could lead to an enhancement of the correlation scale of a putative large-scale magnetic field, as discussed in in [10, 11] (see also [12] and references therein). The limits $R_{\text{kin}} \gg 1$ and $R_{\text{magn}} \gg 1$ are customarily assumed in the scrutiny of hydromagnetic turbulence where both the magnetic flux and the magnetic helicity

²In what follows we shall assume a conformally flat Friedmann-Robertson-Walker background geometry $\bar{g}_{\mu\nu} = a^2(\tau)\eta_{\mu\nu}$ where $\eta_{\mu\nu}$ is the Minkowski metric, τ the conformal time coordinate and $a(\tau)$ the scale factor; the conformal time coordinate τ is related to the cosmic time t as $a(\tau)d\tau = dt$.

are conserved since $\nu_{\text{magn}} \ll 1$ i.e.³

$$\frac{d}{d\tau} \int_{\Sigma} \vec{B} \cdot d\vec{\Sigma} = -\nu_{\text{magn}} \int_{\Sigma} \vec{\nabla} \times (\vec{\nabla} \times \vec{B}) \cdot d\vec{\Sigma}, \quad \frac{d}{d\tau} \int_V d^3x \vec{A} \cdot \vec{B} = -2\nu_{\text{magn}} \int_V d^3x \vec{B} \cdot (\vec{\nabla} \times \vec{B}), \quad (2)$$

where V and Σ are a fiducial volume and a fiducial surface moving with the conducting fluid; \vec{B} and \vec{A} denote the comoving magnetic field and the comoving vector potential. In the ideal hydromagnetic limit (i.e. $\sigma \rightarrow \infty$, $\nu_{\text{magn}} \rightarrow 0$ and $R_{\text{magn}} \rightarrow \infty$) the flux is exactly conserved and the number of links and twists in the magnetic flux lines is also preserved by the time evolution. If $R_{\text{kin}} \gg 1$ and $R_{\text{magn}} \leq \mathcal{O}(1)$ the system is still turbulent; however, since the total time derivative of the magnetic flux and of the magnetic helicity are both $\mathcal{O}(\nu_{\text{magn}})$ the terms at the right hand side of Eq. (2) cannot be neglected. Finally, if $R_{\text{magn}} \gg 1$ and $R_{\text{kin}} \ll 1$ the fluid is not kinetically turbulent but the magnetic flux is conserved.

For $T > \text{Mev}$ the Reynolds numbers can be viewed as a measure the entropy stored in a given Hubble volume. The entropy stored within the Hubble volume $V_H \sim 4\pi H^{-3}(t)/3$ is directly expressible in terms of the Reynolds numbers:

$$S_H = \frac{4}{3}\pi s H^{-3} = \frac{8\pi^3}{135} \frac{N_{\text{eff}}}{v_{\text{rms}}^3 \alpha_{\text{em}}^6} R_{\text{kin}}^3, \quad (3)$$

where s is the entropy density of the plasma. The adiabatic expansion implies that the total entropy over a comoving volume is conserved. Conversely S_H , i.e. the entropy stored in the Hubble volume, increases as the plasma cools down since the temperature redshifts as a^{-1} (where a is the scale factor) but the Hubble volume typically increases faster than a^3 both during radiation (i.e. $V_H \sim a^6$) and during matter (i.e. $V_H \sim a^{9/2}$). Up to numerical factors $S_H \sim \mathcal{O}(10^{64})$ right before electron-positron annihilation and $S_H \sim \mathcal{O}(10^{48})$ in the symmetric phase of the electroweak theory. The maximal entropy stored in the Hubble volume today⁴ is obtained by integrating the entropy density of the Cosmic Microwave Background radiation (CMB in what follows) over the present value of the Hubble volume:

$$S_{\gamma} = \frac{4}{3}\pi s_{\gamma} H_0^{-3} \simeq 1.43 \times 10^{88} \left(\frac{h_0}{0.7} \right)^{-3}, \quad s_{\gamma} = \frac{4}{45} \pi^2 T_{\gamma}^3, \quad (4)$$

where $T_{\gamma} = 2.725$ K. In the standard lore the huge value of the entropy contained in the Hubble volume is the result of an appropriate theory of the initial conditions, since the adiabaticity condition can only be mildly violated after inflation and for a standard thermal history.

According to Eqs. (3)–(4) it would be tempting to establish a causal connection between the Hubble entropy of the CMB and the largeness of the Reynolds numbers. While such a connection cannot be excluded for $T > \text{MeV}$, Eqs. (3) and (4), taken at face value, would imply that the kinetic and magnetic

³In hydromagnetic turbulence it is customarily assumed that $Pr_{\text{magn}} \simeq 1$ while the flow is incompressible (i.e. the bulk velocity of the plasma is solenoidal) [1, 2]. This is not necessarily the case in the early Universe, as we shall see.

⁴We do not consider here the possibility of a gravitational entropy associated with a Hubble screen approximately saturating the Hawking-Bekenstein bound and implying $S_{\text{screen}} \sim H_0^{-2} M_{\text{P}}^2 \simeq \mathcal{O}(10^{120}) \gg S_{\gamma}$.

Reynolds numbers must reach their maximum for $T \ll \text{MeV}$. The latter conclusion is incorrect insofar as the thermal diffusion coefficient sharply increases after electron-positron annihilation while the conductivity is only suppressed as $\sqrt{\bar{T}/(m_e a)}$ where $\bar{T} = aT$ is the comoving temperature and m_e is the electron mass. While prior to electron-positron annihilation $R_{\text{kin}} \gg 1$ and $R_{\text{magn}} \gg 1$, after e^+e^- annihilation R_{magn} is still very large, R_{kin} gets smaller than 1 and Pr_{magn} sharply increases. Indeed, prior to last scattering, the thermal diffusivity is dominated by Thomson scattering and the concentration of the charge carriers is not of the order of the photon concentration (as for $T > \text{MeV}$) but ten orders of magnitude smaller i.e. $n_e = \eta_b n_\gamma$ where n_γ is the comoving concentration of the photons and $\eta_b = \mathcal{O}(10^{-10})$ is the ratio of the baryonic concentration to the photon concentration indirectly probed by big-bang nucleosynthesis and affecting the present abundances of light nuclear elements. The thermal diffusion coefficient is then given by:

$$\nu_{\text{th}}(\tau) = \frac{4}{5} c_{\text{sb}}^2(\tau) \lambda_{e\gamma}(\tau), \quad c_{\text{sb}}(\tau) = \frac{1}{\sqrt{3[1 + R_b(\tau)]}}, \quad (5)$$

where the electron-photon mean free path $\lambda_{e\gamma}$ and the ratio between the baryonic matter density and the photon energy density R_b are defined as:

$$\lambda_{e\gamma} = \frac{a_0}{\tilde{n}_e a(\tau) \sigma_{\gamma e}}, \quad R_b(\tau) = \frac{3}{4} \left(\frac{\omega_{b0}}{\omega_{\gamma 0}} \right) \left(\frac{a}{a_0} \right) = \frac{685.62}{z+1} \left(\frac{\omega_{b0}}{0.02258} \right); \quad (6)$$

the baryon matter density ρ_b is the sum of the matter densities and of the ions and electrons; $\sigma_{\gamma e}$ is the electron-photon cross section; $\omega_{b0} = h_0^2 \Omega_{b0}$ and Ω_{b0} is the critical fraction of baryons. For the sake of simplicity we shall adopt, in the explicit estimates, the following fiducial set of parameters

$$(\Omega_{b0}, \Omega_{c0}, \Omega_{de0}, h_0, n_s, \epsilon_{\text{re}}) \equiv (0.0449, 0.222, 0.734, 0.710, 0.963, 0.088), \quad (7)$$

which are determined from the WMAP 7yr data alone [13] in the light of the vanilla ΛCDM scenario. After electron-positron annihilation the conductivity given by binary collisions can be estimated as

$$\sigma(\tau) = \sigma_1 \frac{\bar{T}}{\alpha_{\text{em}}} \sqrt{\frac{\bar{T}}{m_e a}} \frac{1}{\ln \Lambda_C}, \quad \Lambda_C(\bar{T}) = \frac{3}{2e^3} \left(\frac{\bar{T}^3}{\pi n_e} \right)^{1/2} = 1.105 \times 10^8 \left(\frac{\omega_{b0}}{0.02258} \right)^{-1/2}, \quad (8)$$

where $\sigma_1 = 9/(8\pi\sqrt{3})$ depends on the way multiple scattering is estimated and Λ_C is the argument of the Coulomb logarithm.

Let us finally come to a more detailed estimate of the velocity field prior to last scattering. The bulk velocity is defined as the center of mass velocity of the positive and negative charge carriers present in the globally neutral plasma⁵. The customary assumptions of hydromagnetic turbulence imply a solenoidal bulk velocity field with $Pr_{\text{magn}} \simeq \mathcal{O}(1)$ [1, 2]. Conversely, prior to photon decoupling, the

⁵If the charge carriers coincide with electrons and ions, denoting with m_i and m_e the masses of the electrons and ions the bulk velocity of the plasma is defined as $\vec{v}_b = (m_e \vec{v}_e + m_i \vec{v}_i)/(m_e + m_i)$. The center of mass velocity of the electron-ion system is often called baryon velocity.

bulk velocity of the plasma is not solenoidal (i.e. $\vec{\nabla} \cdot \vec{v}_b \neq 0$) and $Pr_{\text{magn}} \gg 1$. This means that the plasma is compressible and the divergence of the bulk velocity is directly affected by the large-scale curvature fluctuations. Using Eq. (6) and assuming $v_{\text{rms}} = 1$ the kinetic and the magnetic Reynolds numbers can be estimated as $R_{\text{kin}} \simeq 0.03$ and $R_{\text{magn}} \simeq 1.30 \times 10^{20}$ for a typical last-scattering redshift $z \sim 1090$. The upper limit obtained on R_{kin} by assuming that v_{rms} coincides with the speed of light can be made more stringent since the large-scale flow, prior to last scattering, can be determined from the evolution equations of the baryon-photon system.

For $T < \text{MeV}$ and around last-scattering the differences between the velocities of the baryons and of the photons are quickly washed out because of the tight-coupling between ions, electrons and photons. Recalling that $\mathcal{H} = \partial_\tau \ln a = aH$, to lowest order in the tight-coupling approximation the truncated set of hydromagnetic equations reads [14]

$$\partial_\tau \vec{v}_{\gamma b} + \frac{\mathcal{H} R_b}{R_b + 1} \vec{v}_{\gamma b} = \frac{R_b}{R_b + 1} \frac{\vec{J} \times \vec{B}}{\rho_b a^4} - \frac{\vec{\nabla} \delta_\gamma}{4(R_b + 1)} - \vec{\nabla} \phi + \nu_{\text{th}} \nabla^2 \vec{v}_{\gamma b}, \quad (9)$$

$$\partial_\tau \delta_b = 3\partial_\tau \psi - \vec{\nabla} \cdot \vec{v}_{\gamma b} + \frac{\vec{J} \cdot \vec{E}}{\rho_b a^4}, \quad \partial_\tau \delta_\gamma = 4\partial_\tau \psi - \frac{4}{3} \vec{\nabla} \cdot \vec{v}_{\gamma b},$$

$$\partial_\tau \vec{B} = \vec{\nabla} \times (\vec{v}_{\gamma b} \times \vec{B}) + \nu_{\text{magn}} \nabla^2 \vec{B} + \vec{\nabla} \times \left(\frac{\vec{\nabla} p_e}{en_e} \right) - \frac{1}{4\pi en_e} \vec{\nabla} \times [(\vec{\nabla} \times \vec{B}) \times \vec{B}], \quad (10)$$

where δ_b and δ_γ denote the density contrasts of the baryons and of the photons in the longitudinal gauge; ϕ and ψ denote, respectively, the (00) and (ii) fluctuations of the conformally flat background geometry adopted in the present discussion. The total Ohmic current \vec{J} obeys an evolution equation which can be reduced to a consistency condition as in the Eckart approach to relativistic thermodynamics; such a relation is given by:

$$\vec{J} = \sigma \left(\vec{E} + \vec{v}_{\gamma b} \times \vec{B} + \frac{\vec{\nabla} p_e}{en_e} - \frac{\vec{J} \times \vec{B}}{en_e} \right), \quad (11)$$

where $n_e = a^3 \tilde{n}_e$. The system of Eqs. (9)–(11) is supplemented by the evolution equations of the curvature perturbations which have been written elsewhere in the longitudinal gauge and even in full gauge-invariant terms (see [14] and references therein).

In previous studies [3, 14] the emphasis has been to see which are the effects of the large-scale magnetic fields on the scalar modes of the geometry. The hierarchy between R_{kin} and R_{magn} after electron-positron annihilation suggests the possibility of addressing also the complementary part of the problem, i.e. the effect of the large-scale flow on the evolution of the magnetic field. Neglecting the terms which are quadratic in the magnetic field as well as the thermoelectric term (i.e. the last two contributions in Eq. (11)) the large-scale flow can be determined by using a standard WKB analysis giving, in Fourier space,

$$\vec{v}_{\gamma b}(k, \tau) = i \hat{k} \overline{M}_{\mathcal{R}}(k, \tau) \sin[kr_s(\tau)] e^{-k^2/k_d^2}, \quad \hat{k} = \frac{\vec{k}}{k},$$

$$r_s(\tau) = \int_0^\tau c_{\text{sb}}(\tau') d\tau', \quad \frac{1}{k_d^2(\tau)} = \frac{1}{2} \int_0^\tau \nu_{\text{th}}(\tau') d\tau', \quad (12)$$

where $r_s(\tau)$ and $k_d(\tau)$ denote, respectively, the sound horizon and the typical scale of diffusive damping; $\overline{M}_{\mathcal{R}}(k, \tau)$ encodes the normalization inherited from the (adiabatic) curvature perturbations which are the only source of large-scale inhomogeneities in the vanilla Λ CDM scenario:

$$\overline{M}_{\mathcal{R}}(k, \tau) = 3^{1/4} c_{\text{sb}}^{3/2}(\tau) \left[\frac{1}{c_{\text{sb}}^2(\tau)} - 2 \right] T_{\mathcal{R}}(\tau) \mathcal{R}_*(\vec{k}), \quad T_{\mathcal{R}}(\tau) = 1 - \frac{\mathcal{H}}{a^2} \int_0^\tau a^2(\tau') d\tau'. \quad (13)$$

As a result of Eqs. (9)–(10), the velocity field (12) is not solenoidal. This situation differs from standard hydromagnetic turbulence and it is closer to the situation of acoustic turbulence (see e.g. [16], first paper). Following the standard conventions [13], the correlation function of curvature perturbations in Fourier space is:

$$\langle \mathcal{R}_*(\vec{p}) \mathcal{R}_*(\vec{q}) \rangle = \frac{2\pi^2}{q^3} \delta^{(3)}(\vec{q} + \vec{p}) \mathcal{P}_{\mathcal{R}}(q), \quad \mathcal{P}_{\mathcal{R}}(q) = \mathcal{A}_{\mathcal{R}} \left(\frac{q}{q_p} \right)^{n_s-1}, \quad (14)$$

where, according to Eq. (7), $n_s = 0.963$ and $\mathcal{A}_{\mathcal{R}} = (2.43 \pm 0.11) \times 10^{-9}$; $q_p = 0.002 \text{ Mpc}^{-1}$ denotes the pivot scale at which the power spectrum of curvature perturbations is conventionally normalized. Given the intrinsic inhomogeneity of the large-scale velocity flow, it is natural to generalize the kinetic and the magnetic Reynolds numbers of Eq. (1) to Fourier space by keeping the dependence on the wavenumbers in the velocity, the dependence on the redshift in the diffusion coefficients and by choosing $L_v \simeq L_B = 1/k_{\text{phys}}(\tau)$ where $k_{\text{phys}} = k(z+1)$:

$$\begin{aligned} R_{\text{kin}}(k, z) &= \mathcal{N}_{\text{kin}} F(k, z), & Pr_{\text{magn}}(k, z) &= \mathcal{N}_{\text{magn}} G(k, z) \\ \mathcal{N}_{\text{kin}} &= \frac{15\pi}{2 \cdot 3^{3/4}} \frac{\alpha_{\text{em}}^2}{m_e^2 H_0} \eta_b n_\gamma \sqrt{\mathcal{A}_{\mathcal{R}}}, & \mathcal{N}_{\text{magn}} &= \frac{9\pi^2}{20\zeta(3)\alpha_{\text{em}} \ln \Lambda_C(\overline{T})} \frac{\sigma_1}{\eta_b} \left(\frac{m_e}{\overline{T}} \right)^{3/2}, \\ F(k, z) &= H_0 [1 - 2c_{\text{sb}}^2(z)] T_{\mathcal{R}}(z, z_{\text{eq}}) \left(\frac{k}{k_p} \right)^{n_s-1} \frac{\sin[k r_s(z)]}{k c_{\text{sb}}^{5/2}(z)}, & G(k, z) &= \frac{c_{\text{sb}}^2(z)}{(z+1)^{3/2}}, \end{aligned} \quad (15)$$

where $\zeta(3) = 1.202$; introducing the function $y(z, z_{\text{eq}}) = [\sqrt{1 + (z_{\text{eq}} + 1)/(z + 1)} - 1]$ the function $T_{\mathcal{R}}(z, z_{\text{eq}})$ can be expressed as:

$$T_{\mathcal{R}}(z, z_{\text{eq}}) = (z+1) \left\{ 1 - \frac{2}{15} \frac{[y(z, z_{\text{eq}}) + 1][3y^2(z, z_{\text{eq}}) + 15y(z, z_{\text{eq}}) + 20]}{[2 + y(z, z_{\text{eq}})]^3} \right\}. \quad (16)$$

In Fig. 1 the kinetic and the magnetic Reynolds numbers are illustrated for the set of cosmological parameters of Eq. (7). The three curves in each plot correspond to three different wavenumbers. For even larger wavenumbers both quantities are oscillating as it can be argued from Eq. (15).

The hierarchy between the kinetic and the magnetic Reynolds number defines naturally a perturbative scheme where the evolution equations of the magnetic field can be averaged over the large-scale

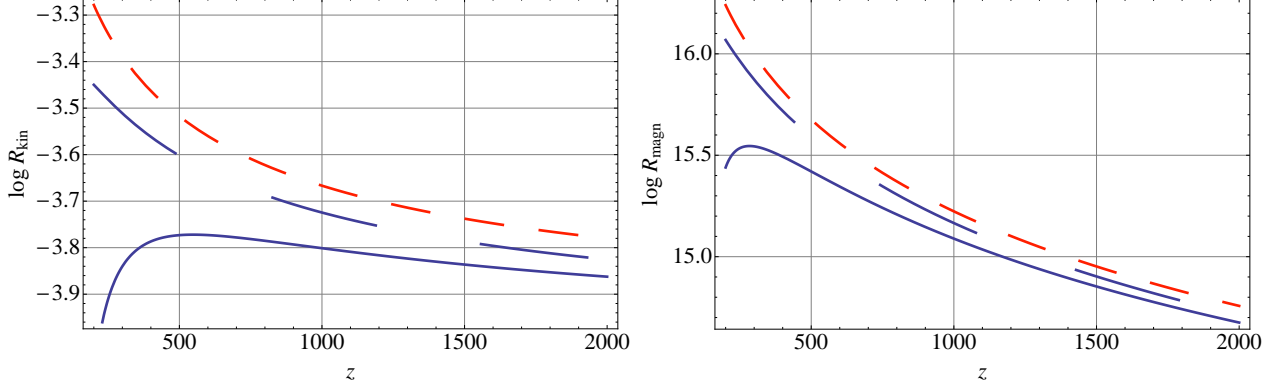


Figure 1: The kinetic Reynolds number (left plot) and the magnetic Reynolds number (right plot) for $k = 0.0002 \text{ Mpc}^{-1}$ (short dashed line), $k = 0.002 \text{ Mpc}^{-1}$ (long dashed line) and $k = 0.004 \text{ Mpc}^{-1}$ (full line). In both plots on the vertical axis the common logarithm of the corresponding quantity is illustrated.

flow. Consider the magnetic diffusivity equation (10) and neglect all the terms which are of higher order in the magnetic field intensity. An iterative solution of Eq. (10) can then be constructed as

$$B_i(\vec{k}, \tau) = \sum_{n=0}^{\infty} B_i^{(n)}(\vec{k}, \tau), \quad \mathcal{G}_k(y) = e^{-k^2 \nu_{\text{magn}} y} \quad (17)$$

$$B_i^{(n+1)}(\vec{k}, \tau) = \frac{(-i)}{(2\pi)^{3/2}} \int_0^\tau \mathcal{G}_k(\tau - \tau_1) d\tau_1 \int d^3q \int d^3p \delta^{(3)}(\vec{k} - \vec{q} - \vec{p}) \\ \times \epsilon_{mni} \epsilon_{abn} (q_m + p_m) v_a(\vec{q}, \tau_1) B_b^{(n)}(\vec{p}, \tau_1), \quad (18)$$

where, for simplicity, ν_{magn} is assumed to be constant in time. From Eq. (18) the first few terms of the recursion are $B_i^{(0)}(\vec{k}, \tau)$, $B_i^{(1)}(\vec{k}, \tau)$ and $B_i^{(2)}(\vec{k}, \tau)$. The term $B_i^{(0)}(\vec{k}, \tau) = \mathcal{G}_k(\tau) B_i(\vec{k})$ where $B_i(\vec{k})$ parametrizes the initial stochastic magnetic field obeying flux conservation (see Eq. (2)):

$$\langle B_i(\vec{k}) B_j(\vec{k}') \rangle = \frac{2\pi^2}{k^3} P_{ij}(\hat{k}) P_B(k) \delta^{(3)}(\vec{k} + \vec{k}'), \quad P_{ij}(\hat{k}) = \delta_{ij} - \hat{k}_i \hat{k}_j. \quad (19)$$

Since the curvature perturbations are distributed as in Eq. (14), the correlation function of the velocity for unequal times can be written as

$$\langle v_i(\vec{q}, \tau) v_j(\vec{p}, \tau') \rangle = \frac{q_i q_j}{q^2} \mathcal{U}(q, |\tau - \tau'|) \delta^{(3)}(\vec{q} + \vec{p}), \quad \mathcal{U}(q, |\tau - \tau'|) = v(q) \tau_c \delta(\tau - \tau'), \quad (20)$$

where, to avoid confusions with vector indices, the subscript γb has been suppressed. The function $v(q)$ appearing in Eq. (20) is

$$v(q) = \tau_c \mathcal{V}(q), \quad \mathcal{V}(q) = \overline{M}_{\mathcal{R}}^2(q, \tau_*) \frac{2\pi^2}{q^3} \mathcal{P}_{\mathcal{R}}(q) \sin^2 [q r_s(\tau_*)] e^{-2q^2/q_d^2}, \quad (21)$$

where τ_* denotes the last-scattering time and the correlation time τ_c is the smallest time-scale when compared with other characteristic times arising in the problem. Because of the exponential suppression of the velocity correlation function for $\tau > \tau_d$ (where τ_d denotes the Silk time [15]), τ_c approximately coincides with τ_d . The form of the correlator given in Eq. (20) is characteristic of Markovian conducting fluids [16, 17].

Denoting with $H_i(\vec{k}, \tau)$ the magnetic field averaged over the fluid flow, the terms containing an odd number of velocities will be zero while the correlators containing an even number of velocities do not vanish i.e. $\langle B_i^{(2n+1)} \rangle = H_i^{(2n+1)} = 0$ and $\langle B_i^{(2n+2)} \rangle = H_i^{(2n+2)} \neq 0$. So, for instance, $\langle B_i^{(1)} \rangle = 0$ while $\langle B_i^{(2)} \rangle = H_i^{(2)}$ is

$$\begin{aligned} H_i^{(2)}(\vec{k}, \tau) &= \frac{(-i)^2}{(2\pi)^3} \int d^3q \int d^3p \int d^3q' \int d^3p' \delta^{(3)}(\vec{k} - \vec{q} - \vec{p}) \delta^{(3)}(\vec{p} - \vec{q}' - \vec{p}') \\ &\times \int_0^\tau d\tau_1 \mathcal{G}_k(\tau - \tau_1) \int_0^{\tau_1} d\tau_2 \mathcal{G}_p(\tau_1 - \tau_2) (q_m + p_m) (q'_{m'} + p'_{m'}) \epsilon_{b m' n'} \epsilon_{a' b' n'} \epsilon_{m n i} \epsilon_{a b n} \\ &\times \langle v_{a'}(\vec{q}', \tau_2) v_a(\vec{q}, \tau_1) \rangle B_{b'}(\vec{p}'). \end{aligned} \quad (22)$$

After averaging the whole series of Eq. (17) term by term the obtained result can be resummed and written as:

$$H_i(\vec{k}, \tau) = \langle B_i^{(0)}(\vec{k}, \tau) \rangle + \langle B_i^{(2)}(\vec{k}, \tau) \rangle + \langle B_i^{(4)}(\vec{k}, \tau) \rangle + \dots = e^{-k^2 \bar{\nu}_{\text{magn}} \tau} B_i(\vec{k}). \quad (23)$$

where the magnetic diffusivity coefficient $\nu_{\text{magn}} = 1/(4\pi\sigma)$ has been renormalized as

$$\bar{\nu}_{\text{magn}} = \nu_{\text{magn}} + v_0, \quad v_0 = \frac{\tau_c}{3} \int \frac{dk}{k} \bar{M}_{\mathcal{R}}^2(k, \tau_*) \mathcal{P}_{\mathcal{R}}(k) \sin^2(k/k_*) e^{-2k^2/k_d^2}, \quad (24)$$

and $k_* = 1/r_s(\tau_*)$. The averaging suggested here has been explored long ago in the related context of acoustic turbulence by Vainshtein and Zeldovich [16] (see also [17]). Prior to decoupling, however, both $\langle v^2 \rangle \propto \mathcal{A}_{\mathcal{R}} \ll 1$ and $R_{\text{kin}} \ll 1$. The resum indicated in Eq. (23) seems then to be more plausible in the present case than in the one of a kinetically turbulent plasma with strong inhomogeneities. If the Markovian approximation is relaxed the velocity correlator becomes

$$\langle v_i(\vec{q}, \tau_1) v_j(\vec{p}, \tau_2) \rangle = \hat{q}_i \hat{q}_j \tilde{v}(q) \Gamma(q, \tau_1, \tau_2) \delta^{(3)}(\vec{q} + \vec{p}), \quad \tilde{v}(q) = \frac{2\pi^2}{q^3} \mathcal{P}_{\mathcal{R}}(q). \quad (25)$$

Assuming for sake of simplicity, $R_b(\tau_*) \ll 1$ and $c_{\text{sb}}(\tau_*) \simeq 1/\sqrt{3}$ we have that

$$\Gamma(q, \tau_1, \tau_2) = \frac{3}{50} \left\{ \cos[q c_{\text{sb}}(\tau_1 - \tau_2)] - \cos[q c_{\text{sb}}(\tau_1 + \tau_2)] \right\} e^{-q^2 \nu_{\text{th}}(\tau_1 + \tau_2)}. \quad (26)$$

The results (24) and (26) lead to physically equivalent estimates and support the conclusion that the diffusivity wavenumber is smaller than expected from the usual arguments of the magnetic diffusivity scale at last scattering. In fact, ignoring the contribution of the bulk velocity in Eq. (10) the diffusivity

wavenumber can be roughly estimated as $k_\sigma \sim \sqrt{\mathcal{H}_* \sigma}$. The standard estimate of k_σ must be compared with the diffusivity scale arising from the effect of the large-scale flow:

$$\begin{aligned} k_\sigma &\simeq 2.55 \times 10^{10} \left(\frac{d_A}{14116 \text{ Mpc}} \right)^{-1/2} \text{ Mpc}^{-1}, \\ k_v &\sim 50 \left(\frac{\mathcal{A}_R}{2.41 \times 10^{-9}} \right)^{-\frac{1}{n_s+1}} \left(\frac{d_A}{14116 \text{ Mpc}} \right)^{-\frac{2}{n_s+1}} \text{ Mpc}^{-1}, \end{aligned} \quad (27)$$

where d_A denotes the (comoving) angular diameter distance to last-scattering for the typical set of fiducial parameters of Eq. (7). The standard analysis based on the magnetic diffusivity scale would imply that all the modes $k > k_\sigma$ are diffused. The presence of large-scale flow implies that diffusion operates already for $k \geq k_v$. This means that the correct diffusion scale to be considered prior to decoupling is not $k_\sigma \sim \mathcal{O}(10^{11}) \text{ Mpc}^{-1}$ but, at most, $k_v \sim \mathcal{O}(50) \text{ Mpc}^{-1}$ which is closer to the Silk damping scale but qualitatively and quantitatively different.

In summary, prior to electron-positron annihilation the large-scale (turbulent) flow can only be determined indirectly from the features of the various phase transitions. After electron-positron annihilation the flow can be inferred directly from the evolution of large-scale curvature perturbations imprinted in the CMB anisotropies. The hierarchy between the kinetic and magnetic Reynolds numbers prior to last scattering suggests an effective description of the evolution of pre-decoupling magnetic which encompasses the conventional approach solely based on the conservation of the magnetic flux. On a more speculative ground the present considerations suggest that the Reynolds numbers are somehow related to the largeness of the Hubble entropy during the early stages of the evolution of the plasma.

References

- [1] D. Biskamp, *Magnetohydrodynamic Turbulence*, (Cambridge University Press, 2003), p. 101.
- [2] H. P. Goedbloed and S. Poedts, *Principles of Magnetohydrodynamics*, (Cambridge University Press, Cambridge 2004), p.71.
- [3] M. Giovannini, Int. J. Mod. Phys. D **13**, 391 (2004).
- [4] M. Giovannini, Phys. Rev. D **61**, 063004 (2000); Phys. Rev. D **61**, 063502 (2000); M. Giovannini and M. Shaposhnikov, Phys. Rev. D **57**, 2186 (1998); Phys. Rev. Lett. **80**, 22 (1998).
- [5] C. F. Von Weizsäcker, Astrophys. J. **114**, 165 (1951); G. Gamow, Phys. Rev. **86**, 251 (1952).
- [6] J. H. Oort, Nature **224**, 1158 (1969); L. M. Ozernoy and A. D. Chernin, Soviet Astron. AJ **12**, 901 (1969) [Astron. Zh. **45**, 1137 (1968)]; L. M. Ozernoy, Soviet Astron. AJ **15**, 923 (1972) [Astron. Zh. **48**, 1160 (1971)].

- [7] P. J. E. Peebles, *Astrophys. and Space Sci.* **11**, 443 (1971).
- [8] J. D. Barrow, *Mon. Not. R. astr. Soc.* **179**, 47 (1977); *Mon. Not. R. astr. Soc.* **178**, 625 (1977).
- [9] C. Hogan, *Phys. Rev. Lett.* **51** 1488 (1983); M. Hindmarsh, A. Everett, *Phys. Rev.* **D58**, 103505 (1998).
- [10] A. Brandenburg, K. Enqvist and P. Olesen *Phys. Rev. D* **54**, 1291 (1996); A. Brandenburg, K. Enqvist and P. Olesen *Phys. Lett. B* **392**, 395 (1997).
- [11] P. Olesen, *Phys. Lett. B* **398**, 321 (1997); P. Olesen, *NATO ASI Series B* **366**, 159 (1998).
- [12] K. Enqvist, *Int. J. Mod. Phys. D* **7**, 331 (1998); D. T. Son, *Phys. Rev. D* **59**, 063008 (1999); M. Christensson, M. Hindmarsh, *Phys. Rev.* **D60**, 063001 (1999);
- [13] C. L. Bennett *et al.*, *Astrophys. J. Suppl.* **192**, 17 (2011); N. Jarosik *et al.*, *Astrophys. J. Suppl.* **192**, 14 (2011); J. L. Weiland *et al.*, *Astrophys. J. Suppl.* **192**, 19 (2011); D. Larson *et al.*, *Astrophys. J. Suppl.* **192**, 16 (2011); B. Gold *et al.*, *Astrophys. J. Suppl.* **192**, 15 (2011); E. Komatsu *et al.*, *Astrophys. J. Suppl.* **192**, 18 (2011).
- [14] M. Giovannini, *Phys. Rev.* **D79**, 121302 (2009); *Phys. Rev.* **D79**, 103007 (2009); M. Giovannini and N. Q. Lan *Phys. Rev.* **D80**, 027302 (2009); M. Giovannini and Z. Rezaei *Phys. Rev.* **D83**, 083519 (2011); M. Giovannini, *Phys. Rev.* **D84**, 063010 (2011).
- [15] J. Silk, *Astrophys. J.* **151**, 459 (1968); P. J. E. Peebles and J. T. Yu, *Astrophys. J.* **162** 815 (1970).
- [16] S. I. Vainshtein, *Sov. Phys. Doklady* **15**, 1090 (1971) [*Dokl. Akad. Nauk SSSR* **195**, 793 (1970)]; S. I. Vainshtein and Ya. B. Zeldovich, *Sov. Phys. Usp.* **15** 159 (1972) [*Usp. Fiz. Nauk.* **106**, 431 (1972)].
- [17] S. I. Vainshtein, *Sov. Phys. JETP* **31**, 87 (1970) [*Zh. Eksp. Teor. Fiz.* **58**, 153 (1970)]; S. I. Vainshtein, *Sov. Phys. JETP* **34**, 327 (1971) [*Zh. Eksp. Teor. Fiz.* **61**, 612 (1971)].